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A List of the Most Prospective Eclipsing Cataclysmic Variables According to the TESS

Gulnur Subebekova [1,*], Makhabbat Adil [1], Serik Khokhlov [1], Aldiyar Agishev [1,2] and Almansur Agishev [1]

[1] Department of Electronics and Astrophysics, Al-Farabi Kazakh National University, Almaty 050040, Kazakhstan; adilmahabbat@gmail.com (M.A.); skhokh88@gmail.com (S.K.); aldiyar.agishev@gmail.com (A.A.); almansur97@gmail.com (A.A.)
[2] Data Science in Astrophysics Research Center, Almaty 050040, Kazakhstan
* Correspondence: gulnursubebekova@gmail.com

Abstract

Eclipsing cataclysmic variables (CVs) are key targets for determining binary system parameters through photometric modeling, yet many of them remain poorly characterized. In this work, we present a list (catalog) of 37 confirmed eclipsing CVs selected based on high-quality and publicly available TESS photometric data. The sample includes both long-period systems (with orbital periods exceeding 4 h), such as Z Cam, U Gem, and nova-like variables, as well as a significant number of SW Sextantis stars. Selection criteria required the presence of clearly defined eclipses and sufficient signal-to-noise ratios for reliable analysis. The catalog provides a foundation for phase-folded light curve studies and future modeling efforts aimed at deriving key physical parameters such as component masses, radii, inclinations, and accretion geometries. Notably, several systems, such as V482 Cam, OZ Dra, ASASSN-14ix, and others, have no previously published physical parameters. Our list is accessible via a dedicated website, where each system will have a separate page, including data from TESS, AAVSO, and ZTF. This resource is intended to support detailed follow-up studies. It may encourage other research groups with observational and modeling expertise to contribute to the investigation of these promising but understudied systems.

Keywords: cataclysmicvariable; eclipsing binaries; nova-like type stars; SW-Sex

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1. Introduction

Cataclysmic variables (CV) are interacting binaries composed of a white dwarf (WD) as the primary and a late-type (K-M type) main-sequence star or a brown dwarf as the secondary [1]. Cataclysmic variable systems with weak or absent magnetic fields in white dwarfs (WD) allow for the formation of a complete accretion disk and are classified as non-magnetic CVs. If the WD has a moderate to strong magnetic field (typically 1–10 MG, or $>10^7$ G), the accretion disk is partially disrupted or completely suppressed. In such magnetic CVs, including intermediate polars (IPs), the accreting material is channeled along magnetic field lines onto the WD's magnetic poles, forming accretion columns or hotspots. In IPs, the WD typically rotates asynchronously with the orbital period, leading to characteristic optical and X-ray pulsations in the spin period of the WD. In nonmagnetic CVs, the secondary star, having filled its Roche lobe, begins to flow through the Lagrangian L_1 point into the white dwarf, forming an accretion disk around the primary. Based on the morphology of light curves, non-magnetic systems are divided into several types:

classical novae (CN), dwarf novae (DN), current novae (RN), and nova-like(NL). Systems that have never been observed for outbursts and with persistently high mass transfer rates (>$10^{-9} M_{\odot} yr^{-1}$) [2] are called nova-like. This mass transfer rate is considered high enough to suppress the disk instability mechanism that causes dwarf-nova-type outbursts. The spectra of nova-like variables resemble those of classical novae (CN) that have settled back to quiescence. Eclipsing nova-like systems have a V-shaped eclipse profile. The eclipse depth varies by 1^m–2^m magnitudes. The well-known deep-eclipse nova-like system is UX UMa.

In this paper, we present a comprehensive and up-to-date list of 37 confirmed eclipsing cataclysmic variables (CVs) with orbital periods ranging from 3 to 6 h. Although several general catalogs of cataclysmic variables have been published, such as the Catalog and Atlas of Cataclysmic Variables [3] and the Catalog of Symbiotic Stars [4] a dedicated catalog focusing exclusively on eclipsing cataclysmic variables (ECVs) has been lacking. The present study addresses this gap by compiling a specialized list that includes both well studied systems and those with limited prior observational coverage.

The systems in our list were selected based on well-defined criteria to ensure the reliability and scientific value of the resulting catalog. Specifically, we included only cataclysmic variables that satisfy the following conditions:

1. Eclipsing nature: The presence of deep and well-resolved eclipses in the light curve, clearly visible in TESS full-frame image (FFI) data.
2. Data quality: Availab ility of high-quality photometric time-series data, primarily from TESS, with a sufficient signal-to-noise ratio (S/N $\geq$ 20) to allow for reliable eclipse profile analysis and phase-folded light curves.
3. Data coverage: Observation in at least one full TESS sector, preferably more, to ensure adequate temporal coverage for orbital period confirmation.
4. Classification: Systems previously classified as cataclysmic variables in the literature (e.g., from the Catalogue of Cataclysmic Binaries, Low–Mass X-ray Binaries and Related Objects [5] or Catalog and Atlas of Cataclysmic Variables [3], with consistent photometric and spectroscopic signatures.

These selection criteria were applied systematically to pre-existing CV catalogs, which were cross-matched with the TESS target list and light curve archive. The resulting systems offer strong potential for detailed time-series and phase-folded light-curve analysis, which is essential for determining their physical properties and evolutionary state.

The primary goal of this work is not only to offer a reliable resource for the astronomical community, but also to motivate further observational and theoretical investigations. By providing well-structured data and highlighting the unique characteristics of each system, we aim to stimulate interest among research groups with access to modeling tools or observational facilities. Tools such as PHOEBE and CVLab, in conjunction with Gaia distance measurements, enable more accurate determination of system parameters including component masses, radii, orbital inclinations, and accretion geometries.

Importantly, the catalog is made available via an interactive website https://astro.kaznu.info/Catalog_Eclipsing.html (accessed on 11 August 2025), where each system will have a dedicated page featuring TESS, AAVSO, and ZTF light curves, as well as key physical and observational information. The website is currently in development and will continue to be expanded to ensure full documentation for all listed systems. This structure is designed to make the information more intuitive and accessible to the research community.

We anticipate that our list will support a series of follow-up studies on individual systems, some of which will require additional spectroscopic or photometric observations. By identifying and highlighting promising targets, we hope that this work will serve as a catalyst for collaborative efforts aimed at advancing our understanding of eclipsing CVs.

Through such contributions, the physical parameters of many systems may be refined or even redefined, providing deeper insight into their nature and variability. Here, we present an introduction to 37 eclipsing cataclysmic variables (CV) with orbital periods ranging from 3 to 6 h. Section 2 presents fundamental details about the catalog. Following that, a brief description of each object is provided.

2. Catalog

The main catalog is presented in Table 1 and on the website: https://astro.kaznu.info/Catalog_Eclipsing.html (accessed on 11 August 2025). In our list, we present 37 eclipsing cataclysmic variables of all types. This table includes the names of the systems, coordinates, orbital periods, and types. Some of them do not have defined orbital periods. The majority of the objects are long-period systems classified as nova-like or SW Sex-type, with orbital periods ranging from 3 to 7 h. Based on observational evidence, it is assumed that a common feature of nova-like systems might be disc outflow areas [6], as well as the possibility of establishing certain statistical relationships from light curves [7]. SW Sextantis stars are a subclass of cataclysmic variables (CVs) that were first characterized by [8]. These systems are typically eclipsing nova-like variables that exhibit high-velocity emission line wings, reaching speeds of up to 4000 km/s. Such velocities are inconsistent with a standard origin of the accretion disc, suggesting more complex accretion dynamics. The light curve profiles of some of these objects are shown in Figure 1. Each object is described below.

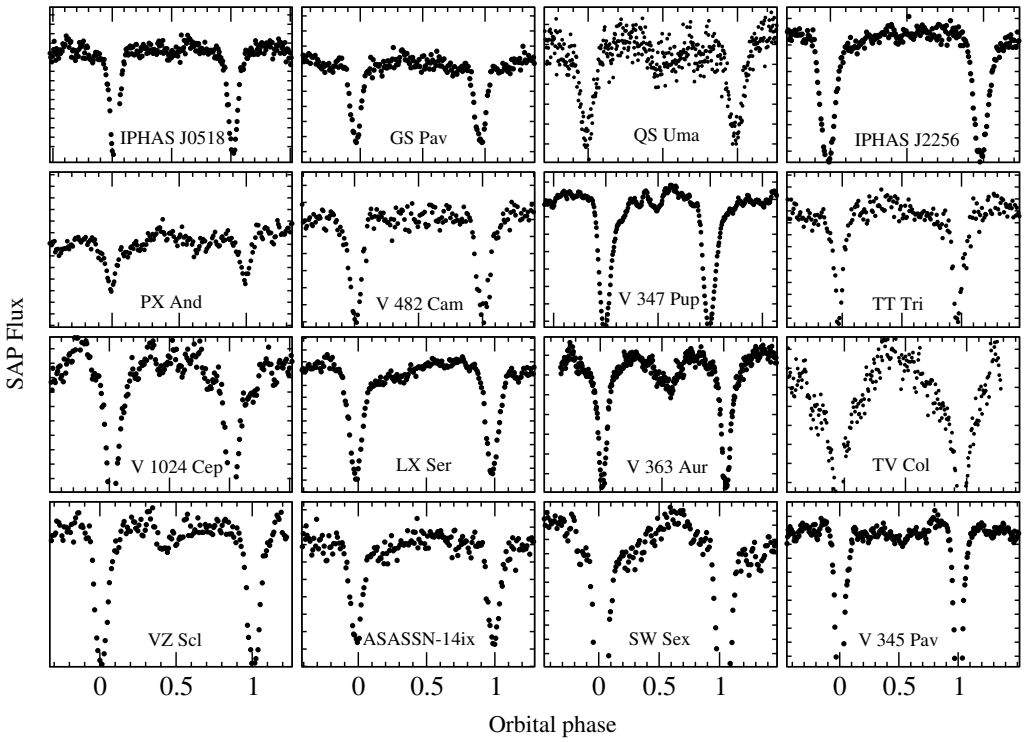


Figure 1. TESS light curves, folded over the orbital periods of 16 objects, all of which exhibit deep eclipses. A secondary minimum is observed in the VZ Scl, V363 Aur, and V347 Pup systems.

2.1. HS 0728+6738

The first light curve of HS 0728+6738 immediately revealed the presence of deep eclipses during which the brightness of the CV drops by ∼2.5 mag. The observation of two consecutive eclipses during the first night provided an estimate of the orbital period of ∼3.2 h [9]. The emission lines are minimally or not at all eclipsed by the secondary star, indicating that their formation occurs outside the Roche lobe of the white dwarf [9].

2.2. SW Sex

The SW Sex system was first identified during the Palomar-Green survey [10]. Early phase-resolved spectroscopic and photometric studies revealed that the primary source of emission lines originates not from the accretion disk itself but from the optically thin region of the disk wind [11]. The orbital period of the system, p = 0.1349384292, was refined later [12], and the presence of a hot spot was prominently observed. The mass transfer rate in this system is notably high, at $1.5 \times 10^{-8} M_{\odot} \mathrm{yr}^{-1}$. In a subsequent study [13], the system was classified as the prototype of SW Sex systems, with the defining characteristics of such systems determined by spectral data analysis. Although the PG 1012-029 system is among the most extensively studied, many of its fundamental parameters remain incompletely constrained.

2.3. HS 0805+3822

HS 0805+3822 was discovered in 2003 by Szkody as a nova-like cataclysmic variable with a short orbital period of 2.4 ± 0.1 h. The time-resolved spectra show deep-absorption lines, especially in the He I and Balmer lines [14]. Later, the fundamental parameters of the system were determined by modeling the spectrum of FUSE, HST, and SDSS using the Tlusty program [15].

2.4. OZ Dra

OZ Dra is a cataclysmic variable (CV) discovered in the Sloan Digital Sky Survey, which shows deep eclipses with an orbital period of 3.28 h. The spectroscopic characteristics of the system suggest that it belongs to the SW Sex class [16]. The power spectrum of the TESS light curve shows a prominent peak in a period similar to [16], along with a much weaker peak that may indicate the presence of a negative superhump [17].

2.5. WX Ari

The initial studies of WX Ari [18] classified WX Ari as a variable similar to the nova of SW Sex type, with no signs of eclipses. However, subsequent R-band photometric observations in 1995 and 1998 revealed the presence of eclipses, indicating a system inclination of $i = 72^{\circ}$ [19].

2.6. V 1315 Aql

V1315 Aql is an eclipsing binary [20]. The system's spectrum exhibits features typical of SW Sex systems, including narrow single-peaked emission lines of Balmer and He [21] .

2.7. TT Tri

TT Tri is an eclipsing binary, as reported by [22]. The orbital period of the system was measured to be 3.35 h by [23]. Its physical parameters have been estimated within a wide range, with the mass ratio between $0.3 \leq q \leq 0.9$ and the orbital inclination ranging from $76^{\circ} \div 79^{\circ}$ [24].

2.8. CM Del

CM Del is an eclipsing cataclysmic variable (CV) discovered by [25]. The system has an orbital period of 0.162 days [26] and consists of a white dwarf with a mass of 0.48 ± 0.15 solar masses and a temperature ranging from 24,000 to 30,000 K. The power spectrum of the TESS data reveals a strong signal in this period [17].

2.9. DW UMa

DW Ursae Majoris is an eclipsing deep ($V \sim 17.6$) nova-like cataclysmic variable (CV) [27]. It is also a key member of the SW Sextantis subclass, an informal group of

CVs defined by observational characteristics [8]. These systems typically share many, but not all, of the following traits: a high inclination, a very high accretion rate, V-shaped eclipses, orbital periods between 3–4 h, positive or negative superhumps, and single-peaked emission lines that are significantly out of phase with the true dynamical motions of the binary components. The orbital period of DW UMa is 3.27 h, which is characteristic of the SW Sextantis class [28].

2.10. ASASSN-14ix

ASAS-SN is an eclipsing dwarf nova system located beyond the period gap [29]. In study [29], a power spectrum was derived using TESS data. The light curve exhibits outbursts typical of dwarf novae, or it could represent a nova-like system with a hot accretion disk. No other research on this object is currently available, leaving its origin uncertain. We include it in our list of systems for which we will determine the physical parameters.

2.11. VZ Scl

VZ Scl is a nova-like cataclysmic variable with an orbital period of 3.28 h [30]. Its brightness outside of eclipses is measured at 15.6 mag [31]. The spectrum of VZ Scl exhibits broad, double-peaked Balmer and helium lines [31]. Currently, no data are available on the physical parameters of this object.

2.12. PX And (PG 0027+260)

PX And = PG 0027+160 was discovered during the Palomar Green Survey [10]. This system was identified as belonging to the SW Sex group based on its spectrum, which exhibits single-peaked emission lines [8]. A periodogram analysis based on 2000 observations revealed three signals with periods of 0.142, 4.8, and 0.207. The first two signals suggest that the system contains negative superhumps [32].

2.13. HS 0455+8315 (V 1024 Cep)

V1024 Cep (known initially as HS 0455+0831) was identified as an eclipsing SW Sex-type star in the Hamburg Quasar Survey. The system has an orbital period of 3.57 h [23]. The first trailing spectra of the H emission line indicate the presence of a high-speed component that is not associated with the primary or secondary component [33].

2.14. HS 0220+0603

HS 0220+0603 is an eclipsing cataclysmic variable of the SW Sex type with an orbital period of 3.58 h, initially identified in the Hamburg Quasar Survey [34]. Spectral studies conducted during the low state of the system provided temperature estimates for its components: $T_{eff} = 30.000$ K for the WD and $T_{eff} = 3.200$ K for the donor star [34].

2.15. GS Pav

GS Pav was initially discovered and designated as the star S 7040 [35]. It was initially classified as a dwarf nova-type cataclysmic variable [3]. Subsequent photometric studies determined an orbital period of 3.72 h and, based on its light curve, the system was later reclassified as a CV similar to a nova-like class [36].

2.16. BH Lyn

BH Lyn was identified as an ultraviolet (UV) source with a magnitude of $V = 15.58$ and classified as CV [10]. Ref. [37] determined its orbital period to be 3.74 h. Spectroscopic studies of PG 0818+513 confirmed this orbital period [38]. The presence of negative superhumps was detected with a period of 0.145 days [39].

2.17. LX Ser

LX Ser was initially identified as a cataclysmic variable star [40], with an orbital period of 3.48 h [41]. The orbital inclination and mass ratio were determined to be $i = 79^\circ$ and $q = 0.5$, respectively [42]. However, no new parameter information has been reported.

2.18. NS Cnc (SDSS J081256.85+191157.8)

NS Cnc was identified as an eclipsing nova-like system by the Sloan Digital Sky Survey [43]. Its orbital period was determined to be 0.1601517(79) days by [44]. Another study detected negative superhumps with a period of 0.148159 days [45].

2.19. UU Aqr

UU Aqr is an eclipsing variable star [46] with an orbital period of 3.56 h. Its spectrum features single-peaked emission lines, which are characteristic of SW Sextantis systems [47]. Observations by [47] provided parameter estimates: an inclination of $i = 78^\circ$ and a mass ratio of $q = 0.3$.

2.20. HBHA 4705-03

HBHA 4705-03 is an eclipsing binary system discovered by [48]. The only available spectroscopic and photometric studies provide information on the component masses ($M_{WD} = 0.54 \pm 0.1 M_\odot$) and the orbital inclination of the system ($i = 71.8^\circ \pm 0.7$).

2.21. HS 2325+8205

HS 2325+8205 is an eclipsing cataclysmic variable (CV) identified in the Hamburg Quasar Survey [49]. The system experienced 44 outbursts between 2007 and 2009 [50]. Later observations revealed outbursts with a duration of 12–14 days, leading to its classification as a CV of Z Cam type [51]. However, no physical parameters of the system have been estimated.

2.22. V 416 Dra

V416 Dra is a poorly studied eclipsing cataclysmic variable (CV), identified by [52] based on optical spectra from the Sloan Digital Sky Survey. During an outburst, the radius of the accretion disk increases, while in quiescence a total eclipse of the disk occurs [53].

2.23. V 539 Vel (2MASS J09440940-5617117)

V539 Vel is an eclipsing binary identified as a cataclysmic variable by [54]. The orbital period was estimated by [55] as 0.1879340 days. The physical parameters of the system were determined by [55], with a mass ratio of $q = 0.66 \div 0.83$ and an orbital inclination of $i = 84^\circ \div 90^\circ$.

2.24. DQ Her

DQ Her is the only eclipsing cataclysmic variable of the classical nova type in our catalog. Its brightest outburst was recorded in 1934, reaching a magnitude of 1.3 mag, while outside the outburst it has a magnitude of 14.3 mag [56]. Recent X-ray studies suggest the presence of a possible magnetic jet [57]. BVRI photometry reveals deep eclipses [58]. Although the system has been extensively studied, its parameters have yet to be determined.

2.25. UX UMa

UX UMa is a prototype of a nova-like cataclysmic variable with deep eclipse. The orbital period of the object is 4.7 h [59]. The proposed disk wind model for the system is given in [60]. Refs. [60,61] estimated the system parameters as $q = 1$, $M_1 = 0.47 \pm 0.07 M_\odot$, $M_2 = 0.47 \pm 0.10 M_\odot$, and $i = 71^\circ$. The initial parameters for UX

UMa include $M_1 = 0.7M_\odot$, $M_2 = 0.47M_\odot$, and the mass transfer rate was estimated to be $\dot{M} = 6.6 \times 10^{17}$ g/s [62]. As observed, the system parameters vary over a wide range.

2.26. V 345 Pav

V345 Pav is an eclipsing binary discovered by the Edinburgh Cape Blue Object Survey, with an orbital period of 4.75 h [63]. Its spectrum displays typical nova-like emission lines, and the system is relatively bright, with an out-of-eclipse magnitude of around 13.7 mag. V345 Pav remains one of the less extensively studied eclipsing cataclysmic variables.

2.27. FO Aqr

FO Aqr is an eclipsing intermediate polar cataclysmic variable (CV) [64] that exhibits two distinct luminosity states: a normal state with an average visual magnitude of $V \approx 13.5$ mag, and a low state around $V \approx 15.5$ mag [65]. The system has an orbital period of 4.85 h and an estimated orbital inclination of approximately $i = 70^\circ$ [66]. Previous studies suggest a wide range of values for the mass of the white dwarf, varying between $M_1 = 0.45M_\odot$ and $M_1 = 1.22M_\odot$ [67–70]. The mass of the secondary star is estimated to lie between $M_2 = 0.4M_\odot$ and $M_2 = 0.46M_\odot$ [70,71]. Given the variability in parameter estimates, further observational studies are necessary to determine the system's physical characteristics accurately.

2.28. IPHAS J051814.34+294113.2

IPHAS J051814 is an eclipsing cataclysmic variable identified through the Isaac Newton Telescope Photometric H_α Survey (IPHAS) [72]. Photometric analysis using data from TESS and ZTF revealed two types of outbursts with average recurrence intervals of approximately 20.7 and 12.9 days, respectively. The system is classified as a Z Cam-type dwarf nova with an orbital period of 4.95 h [73]. Due to its eclipsing nature and regular variability, IPHAS J051814 presents an excellent candidate for detailed modeling. This system is included in our target sample, for which system parameters will be determined using our dedicated computational tool.

2.29. TV Col

TV Col is an intermediate polar cataclysmic variable initially identified as an X-ray source [74]. The system has an orbital period of approximately 5.5 h [75]. In January 2001, it exhibited a short-duration outburst that lasted about six days [76].

2.30. IPHAS J225608.45+595430.0 (2MASSJ22560844+5954299)

IPHAS J225608.45+595430.0 was identified as a cataclysmic variable in 2008 [77]. Photometric observations conducted between 2013 and 2014 revealed light curve characteristics typical of nova-like cataclysmic variables. The estimated temperatures of the components are $T_{WD} = 22,000$ K for the primary and the M-type secondary. The system parameters include a mass ratio of $q = 1.004 \pm 0.0013$ and an orbital inclination of approximately $i = 78.8 \pm 0.1^\circ$ [78].

2.31. RW Tri

RW Tri is a well-known, deeply eclipsing, nova-like cataclysmic variable. The orbital period of the system has been estimated to be approximately 5.6 h [79] and has previously determined the physical parameters of RW Tri [7]. In this work, our objective was to analyze the temporal evolution of the system parameters using photometric data from TESS.

2.32. V 347 Pup

V347 Pup has been classified as a nova-like cataclysmic variable (NL CV) [80]. The orbital period of the system was determined to be 5.56 h. The masses of the components were estimated as $M_1 = 0.73 \pm 0.05 M_\odot$ and $M_2 = 0.54 \pm 0.06 M_\odot$, resulting in a mass ratio of $q = 0.73 \pm 0.05$. The orbital inclination was found to be $i = 85 \pm 2.1^\circ$ [81]. So far, no updated system parameters have been estimated.

2.33. DO Leo

DO Leo is a long-period eclipsing nova-like variable. It was first discovered during a survey for faint blue stars at high galactic latitudes, with an apparent magnitude of $V \sim 16$ [82]. The orbital period of the system was estimated to be 5.62 h. Its spectrum exhibits single-peaked emission lines, which is unusual for systems with high orbital inclination [83]. The system parameters have not been determined previously.

2.34. AC Cnc

AC Cnc is an eclipsing cataclysmic variable with a long period (7.2 h), originally discovered by [84]. Based on previous studies [85], it was classified as a nova-like system. Spectroscopic analysis yielded component masses of $M_1 = 0.82 \pm 0.13 M_\odot$ and $M_2 = 1.02 \pm 0.14 M_\odot$ [86]. However, other studies have estimated the primary and secondary masses to be $M_1 = 0.76 M_\odot$ and $M_2 = 0.77 M_\odot$, respectively [87]. This significant discrepancy motivates a reevaluation of the system parameters using the updated distance provided by GAIA.

2.35. V 363 Aur

V363 Aur is a long-period (7.42 h) eclipsing cataclysmic variable, discovered by [88]. The initial light curve modeling produced the system parameters of the mass ratio $q = 1$ and the orbital inclination $i = 72.7^\circ$ [89]. Later, analysis of the radial velocity curves of the He II emission line provided revised estimates: an inclination of $i = 70^\circ$, mass ratio $q = 1.12 \pm 0.04$, with component masses $M_1 = 0.86 \pm 0.08 M_\odot$ and $M_2 = 0.77 \pm 0.04 M_\odot$ [90]. These discrepancies indicate that the system parameters of this object remain uncertain.

2.36. V 902 Mon

V902 Mon is a deeply eclipsing intermediate polar cataclysmic variable, classified by [91]. The orbital period of the system has been estimated to be 8.16 h [91]. Photometric data from both AAVSO and TESS indicate that the system exhibits characteristics of a variable intermediate polar accreting overflow [92].

2.37. QS UMa

This object has not been studied in detail before. Only its coordinates and stellar magnitude, as reported by Gaia, are currently known.

Table 1. Eclipsing cataclysmic variables.

Name	RA J2000,0	DEC J2000,0	Mag G	P_{orb}	Type	Ref.
HS 0728+6738	07 33 41.4	+67 32 15	16.03	3.21 h	SW Sex	[9]
SW Sex	10 15 09.3	−03 08 32	14.37	3.24 h	SW Sex	[10]
HS 0805+3822	08 09 08.3	+38 14 06	15.24	3.2 h	SW Sex	[14]
OZ Dra	13 27 23.3	+65 28 54	18.34	3.28 h	SW Sex	[16]
WX Ari	02 47 36.2	+10 35 37	17.81	3.34 h	SW Sex	[18]
V1315 Aql	19 13 54.5	+12 18 03	14.30	3.35 h	SW Sex	[20]
TT Tri	01 31 59.8	+29 49 22	15.26	3.36 h	SW Sex	[22]

Table 1. *Cont.*

Name	RA J2000,0	DEC J2000,0	Mag G	P_{orb}	Type	Ref.
CM Del	20 24 56.9	+17 17 54	13.88	3.38 h	SW Sex	[25]
DW UMa	10 33 52.8	+58 46 54	14.25	3.27 h	SW Sex	[27]
ASASSN-14ix	22 55 04.1	−34 49 17	16.84	3.46 h	-	[29]
VZ Scl	23 50 09.2	−26 22 52	15.43	3.47 h	-	[30]
PX And	00 30 05.8	+26 17 26	14.85	3.51 h	-	[10]
HS 0455+8315	05 06 48.2	+83 19 23	15.14	3.57 h	SW Sex	[23]
HS 0220+0603	02 23 01.6	+06 16 49	16.89	3.58 h	-	[34]
GS Pav	20 08 07.6	−69 48 58	16.21	3.72 h	-	[35]
BH Lyn	08 22 36.0	+51 05 24	15.17	3.74 h	-	[10]
LX Ser	15 38 00.0	+18 52 03	14.89	3.48 h	-	[40]
NS Cnc	08 12 56.8	+19 11 57	16.08	3.84 h	-	[43]
UU Aqr	22 09 05.7	−03 46 17	13.51	3.92 h	SW Sex	[46]
HBHA 4705-03	22 16 50.3	+46 46 41	14.92	4.12 h	-	[48]
HS 2325+8205	23 26 50.2	+82 22 11	16.22	4.45 h	Z Cam	[49]
V416 Dra	18 57 20.3	+71 31 18	16.23	4.5 h	U Gem	[52]
V539 Vel	09 44 09.3	−56 17 11	16.81	4.51 h	-	[54]
DQ Her	18 07 30.2	+45 51 32	14.59	4.64 h	-	[56]
UX UMa	13 36 40.9	+51 54 49	12.93	4.72 h	-	[59]
V345 Pav	19 35 42.9	−59 08 22	13.42	4.75 h	-	[63]
FO Aqr	22 17 55.3	−08 21 03	13.86	4.84 h	-	[64]
IPHAS J051814.34+294113.2	05 18 14.3	+29 41 13	15.81	4.95 h	Z Cam	[72]
TV Col	05 29 25.5	−32 49 03	13.98	5.48 h	-	[74]
IPHAS J225608.45+595430.0	22 56 08.4	+59 54 30	14.50	5.48 h	-	[77]
RW Tri	02 25 36.1	+28 05 50	13.25	5.56 h	-	[79]
V347 Pup	06 10 33.6	−48 44 25	13.30	5.56 h	-	[80]
DO Leo	10 40 51.2	+15 11 33	16.87	5.6 h	-	[82]
AC Cnc	08 44 27.1	+12 52 31	14.14	7.2 h	Z Cam	[84]
V363 Aur	05 33 33.4	+36 59 32	14.12	7.7 h	SW Sex	[88]
V902 Mon	06 27 46.4	+01 48 11	16.79	8.16 h	-	[91]
QS UMa	09 32 14.8	+49 50 54	16.47	10.04 h	-	

The coordinates of each object were taken from https://gaia.ari.uni-heidelberg.de/singlesource.html (accessed on 11 August 2025), the stellar magnitudes and orbital periods were obtained from the references cited adjacent to each system.

3. Conclusions

We have developed a list (catalog) of cataclysmic eclipsing systems based on TESS light curves. It includes 37 eclipsing systems, half of which are classified as SW Sex systems. The remaining systems are long-period systems with orbital periods greater than 4 h, including Z Cam, U Gem, and nova-like systems. The advantage of studying these systems lies in the fact that many of them have remained poorly investigated. In particular, for several systems, no physical parameters have been reported to date. This includes V482 Cam, OZ Dra, ASASSN-14ix, HS 0455+8315, HBHA 4705-03, HS 2325+8205, V416 Dra, V539 Vel, IPHAS J2256+5953, V902 Mon, and QS UMa. Moreover, for systems with previously determined parameters, an update is required using the latest distance measurements from the GAIA mission. This is particularly important for objects such as WX Ari, PX And, DW UMa, and V1315 Aql, as existing estimates were obtained relatively long ago and need to be refined. Each object in the catalog is planned for detailed investigation, similar to the comprehensive studies already performed for RW Tri [7].

Our analysis underscores the importance of eclipsing systems as powerful laboratories for investigating binary parameters, accretion physics, and evolutionary processes in CVs. The inclusion of both well-known and previously understudied objects expands the scope for detailed photometric and spectroscopic follow-up observations. The interactive catalog website enhances the usability of our dataset, enabling researchers to access and interpret key observational properties with ease.

We hope that this work will encourage further modeling and observational efforts, especially targeting systems with unexplored or poorly constrained parameters. Ultimately, this list of eclipsing binaries serves as both a scientific resource and a springboard for new

investigations that may deepen our understanding of accretion dynamics, binary evolution, and the role of magnetic fields in shaping the diversity of CV behavior.

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Abbreviations

The following abbreviations are used in this manuscript:

CV Cataclysmic Variable

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