We are very grateful to reviewer for valuable and useful comments which improved our paper. Our answers
are:

1. The theoretical resolving power of the high-resolution echelle spectrometer for the Keck Ten-
Meter Telescope is of the order of > 250,000. However, practical realizations may be ~ 36,000. At
wavelengths, λ, of 500 and 4000 Å, the resolution, Δλ, would be 0.014 and 0.11 Å, respectively. It
will be helpful to include a remark regarding resolving power needed to measure reasonably well
the indicated Stark widths 80 kK.

Our answer:
It is worth to add that the theoretical resolving power of the high-resolution echelle spectrometer for the Keck Ten -
Meter Telescope is of the order of $>$ 250,000. However, practical realizations may be approximately 36,000. Resolution needed
for Stark widths shown in Table 1 may be divided in two groups. For lines between 4217.3 \AA \- and 2987.7 \AA \-
needed resolutions are 6,139 - 12,990 for 10,000 K and 16,035 - 35,953 for 80,000 K and. For lines between
1042.2 \AA \- and 505.4 \AA \-
needed resolutions are 61,060 - 292,321 for 10,000 K and 161,090 - 824,685 for 80,000 K. It is important that for Sn IV 4217.3 \AA \- line,
which is in the visible, needed resolution is 6,139 at 10,000 K and 16,035 at 80,000 K,
so that the influence of Stark broadening can be observed with large terrestrial telescopes.

2. It will also help to include (perhaps as a footnote) in Table 1, or near line 30 on page 2, values for
the Doppler broadening at, say 80 kK, for comparison with the predicted values

Our answer:
At the end of paragraph, added as the answer to previous remark, is added:

For this
line, thermal Doppler width is 0.0277 \AA \- and Stark 0.687 \AA \- for T = 10,000 K. The corresponding values for 80,000 K are 0.0784 \AA \-
and 0.263 \AA \- respectively.

3. Line 36, page 2: The classification may be DZ (or if unclear DX); and line 36, page 2, A-type stars
appear to show temperatures <12 kK, so, please edit and use suggested nomenclature.

Our answer:
The title of Ref. 3 is: Discovery of germanium, arsenic, selenium, tin, tellurium and iodine in the atmospheres of cool DO white dwarfs. The title of Ref 4 is: Abundance of Elements beyond the Iron Group in Cool DO White Dwarfs.DO white dwarf stars have helium-rich atmospheres with effective temperatures ranging from 120,000 K to 45,000 K.The temperatures of cool DO white dwarfs mark the hot end of the so-called DB gap, which corresponds to an interval in effective temperatures from 45,000 K to 30,000 K. Since
DZ white dwarfs may have even temperatures lower than 9,000 K it would not be adequate to name cool DO white dwarfs as DZ white dwarfs. In order to be more clear we modified text so that:
"DO white dwarfs \cite{Ch05a, Ch05b}."
is replaced by:
"cool DO white dwarfs \cite{Ch05a, Ch05b}. They mark the hot end of the so-called DB gap, which corresponds to an interval in effective temperatures from 45,000 K to 30,000 K \cite{Ch05b}."

4. There are just a few edits/grammar, e.g., lines 121 on page 5: “, i.e.,” and next line 121 “Contrary,”
or line 127, “corresponding”.

Corrected. Thank you

5. On page 5, line 126, comments RE accuracy of Stark widths: I understand that this is ± 11% or
±30 %, how does this affect the needed resolution for, say, the echelle in point 1.

Our answer:
Maybe we were not clear enough. These are not error bars but differences of Stark widths of homologous transitions in spectra of homologous ions Se IV and Te IV, presented in order to see how regular is the behaviour of Stark widths within this homologous sequence. In order to be clearer we replaced "they are only for" with "widths are only for".